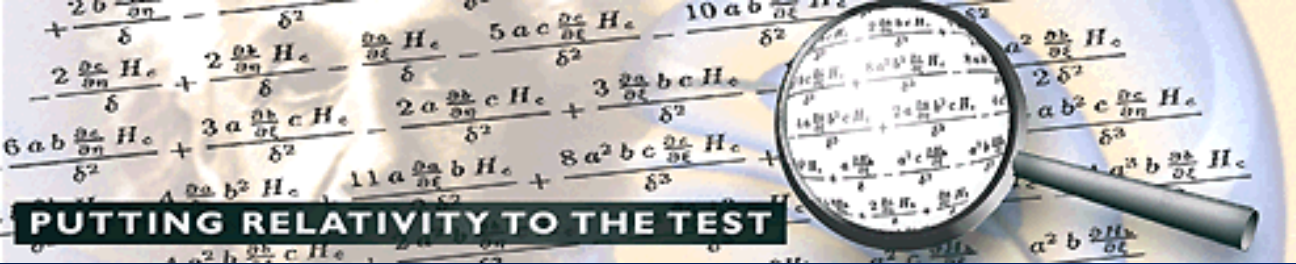


Constraints on time variation of G_N

from Geometrical Test and Perturbation

Seokcheon (sky) Lee 李碩天
Institute of Physics, Academia Sinica
May. 10th. 2012
CYCU HEP & QIS seminar

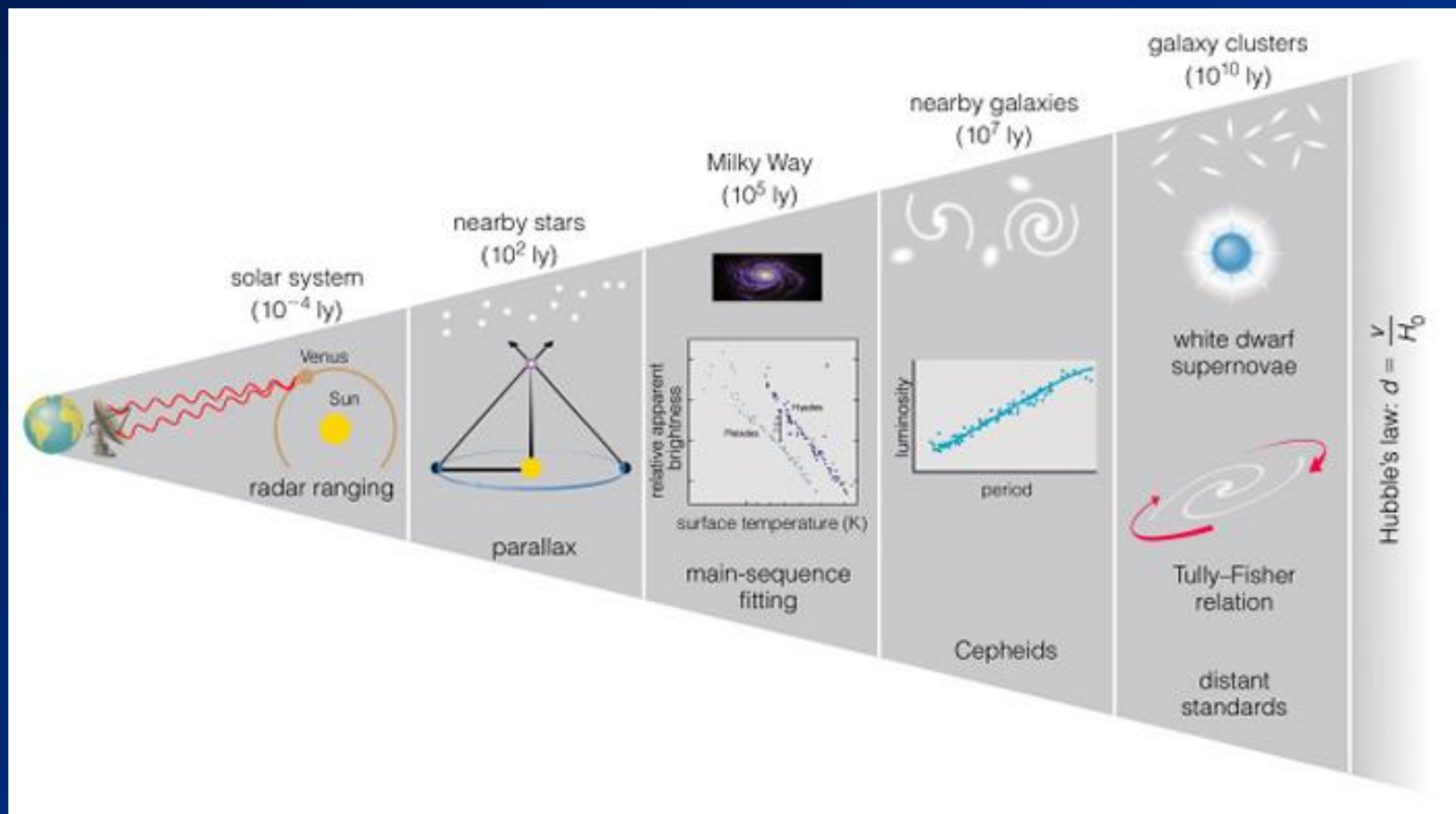
Outline



- 1 Age and Distance.....•
- 2 Observations (SNe, Growth).....•
- 3 Constraints on G_N•
- 4 Conclusion.....•

Cosmic Distance Ladder

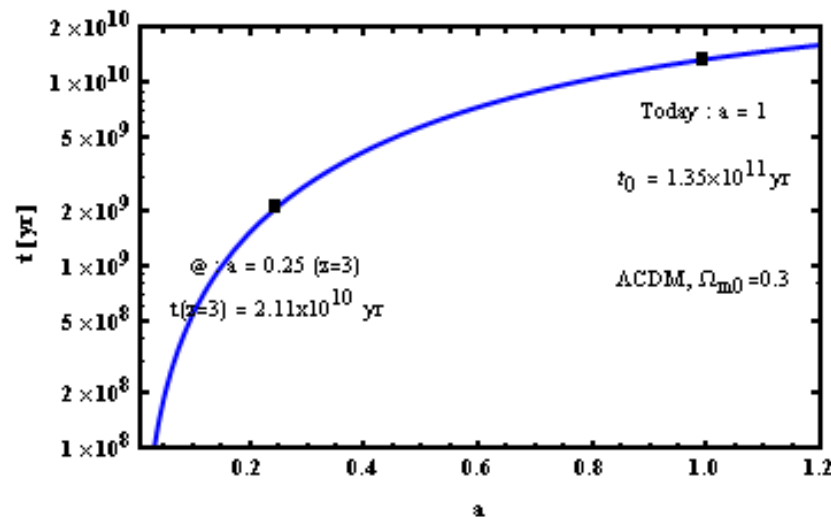
❖ 1 parsec = 3.26 ly



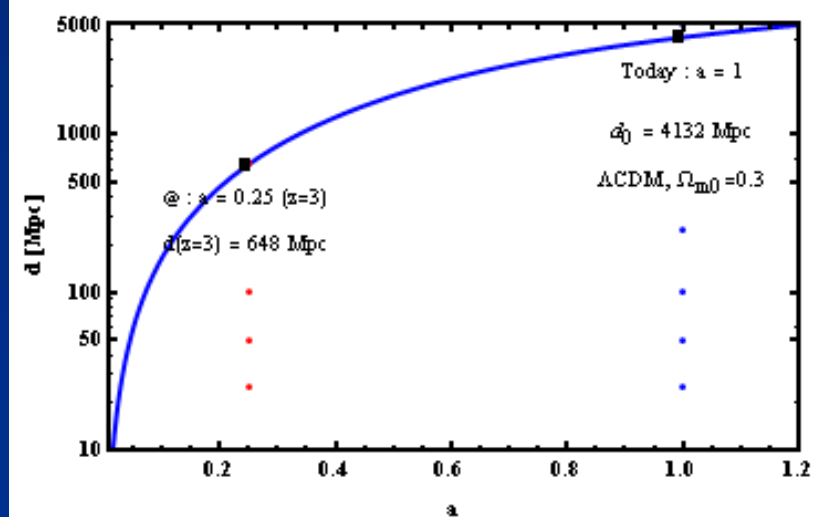
Age and Distance

Need DE?
 t_0 and d_0

t_0



d_0



Discriminate DE from MG

Background can be same

$$\begin{aligned}
 H^2 - \delta H &= \frac{8\pi G_*}{3} \rho_m \rightarrow \delta H \equiv \frac{8\pi G_*}{3} \rho_X \\
 H^2 &\equiv \frac{8\pi G_*}{3} (\rho_m + \rho_X) \\
 \dot{H} &= -4\pi G_* \rho_m - 4\pi G_* (1 + \omega_X) \rho_X \\
 -\Omega_X \omega_X &= 1 + \frac{2}{3} \frac{\dot{H}}{H^2}
 \end{aligned}$$

MG vs DE

Subhorizon scale

Matter Growth

$$\begin{aligned}
 \frac{d \ln \delta_m}{d \ln a} &\equiv \Omega_m^\gamma \\
 \gamma &\equiv \gamma_0 + \gamma_a (1 - a) \\
 \delta_m'' + \left(2 + \frac{H'}{H}\right) \delta_m' - \frac{4\pi G_{eff} \rho_m}{H^2} \delta_m &= 0 \\
 \text{where } \delta_m' &\equiv \frac{d \delta_m}{d \ln a}
 \end{aligned}$$

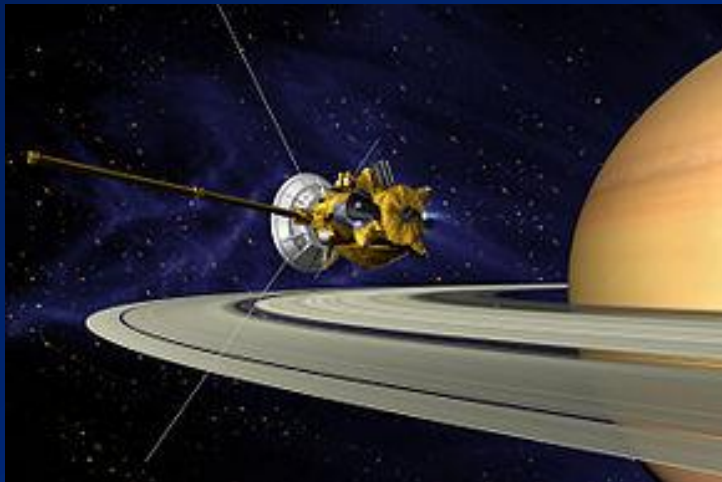
Matter growth different due to G_{eff}

Theoretical Prediction

Quantities	$\omega = \omega_0 + \omega_a(1-a)$		γ
	ω_0	ω_a	
Λ	-1	0	0.56
Quintessence	~ -1	$\neq 0$	~ 0.56
DGP	-0.78	0.32	11/16 (0.69)
f(R)	-0.7	positive	0.43 -0.18 (1-a)
Scalar-Tensor-Gravity	flexible	flexible	Determined from ω

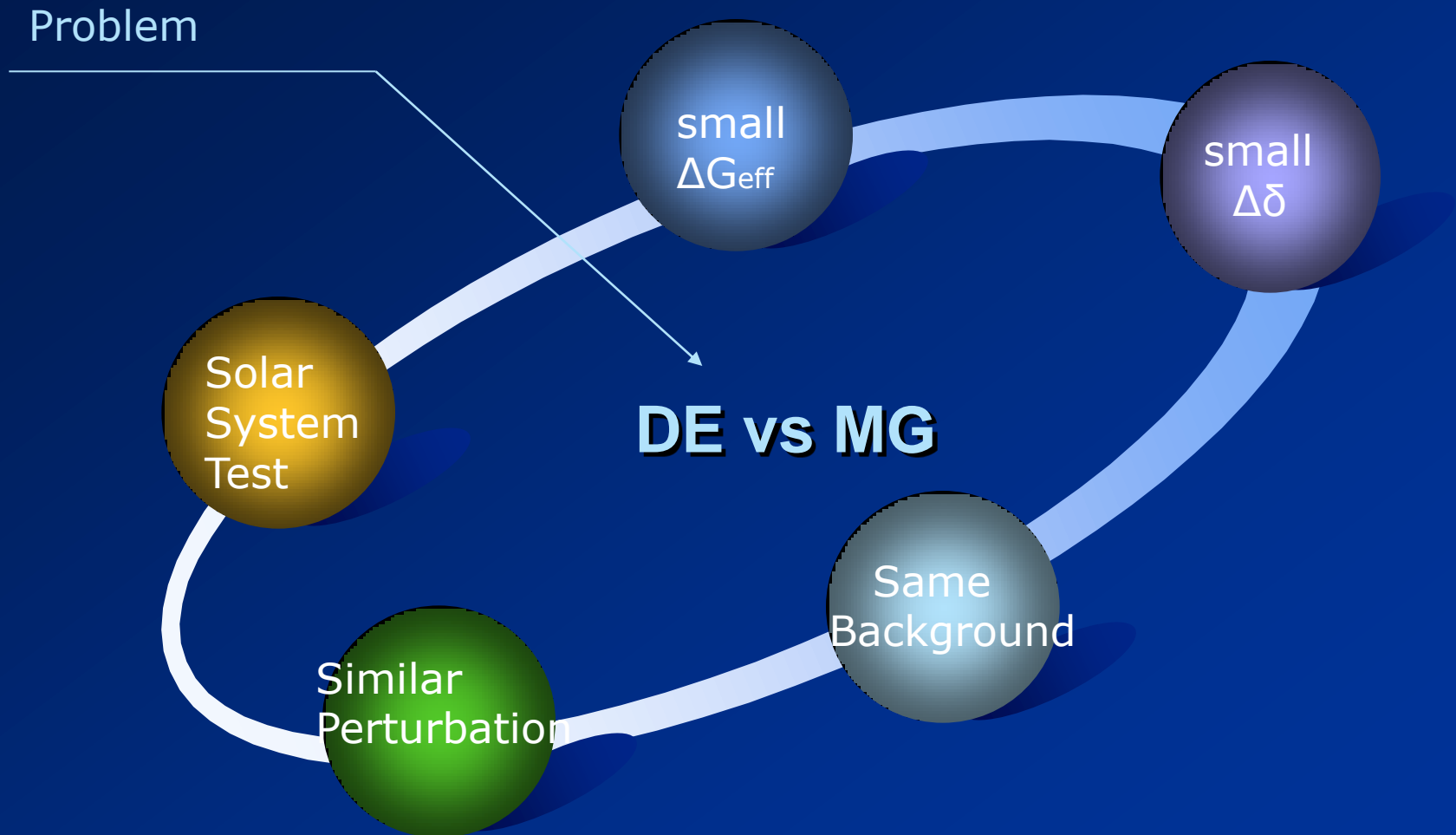
Cassini–Huygens Mission

- ❖ A joint NASA/ESA/ASI spacecraft mission studying the planet Saturn



$$\begin{aligned}\dot{G}/G &= (-0.6 \pm 1.6) \times 10^{-12} \text{ yr}^{-1} \\ &= (-0.009 \pm 0.022) H_0,\end{aligned}$$

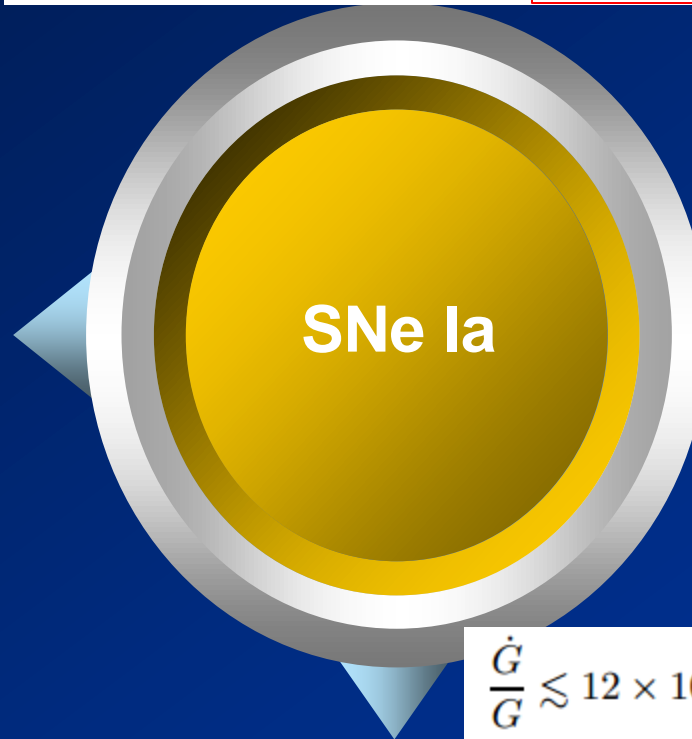
Problem in Discriminating DE from MG



SNe Ia : Standard Candle

(Gaztanaga et.al 2001)

$$m(z) = M_0 + 5 \log d_L + 25 + \frac{15}{4(1+\omega)} \log(1+z)$$



$$M_C \simeq \frac{3.1}{m'} \left(\frac{\hbar c}{G} \right)^{3/2}$$

$$L_p \sim G^{-\gamma}$$

$$M - M_0 = \frac{15}{4} \log \left(\frac{G}{G_0} \right)$$

$$G(z) \equiv G_0 (1+z)^{\frac{1}{1+\omega(z)}}$$

$$\frac{G}{G_0} \lesssim 1.08 \quad ; \quad \Omega_\Lambda \simeq 0.8 \quad , \quad \Omega_M \simeq 0.2$$

$$\frac{\dot{G}}{G} \simeq \left(1 - \frac{G_0}{G} \right) (\Delta t)^{-1}$$

$$\frac{\dot{G}}{G} \lesssim 12 \times 10^{-12} \, h_{70}/\text{yr} \quad ; \quad \Omega_\Lambda \simeq 0.8 \quad , \quad \Omega_M \simeq 0.2$$

SNe Ia : explode when WD reaches Chandrasekhar Mass (M_C)

Relation btw CM and luminosity (L)

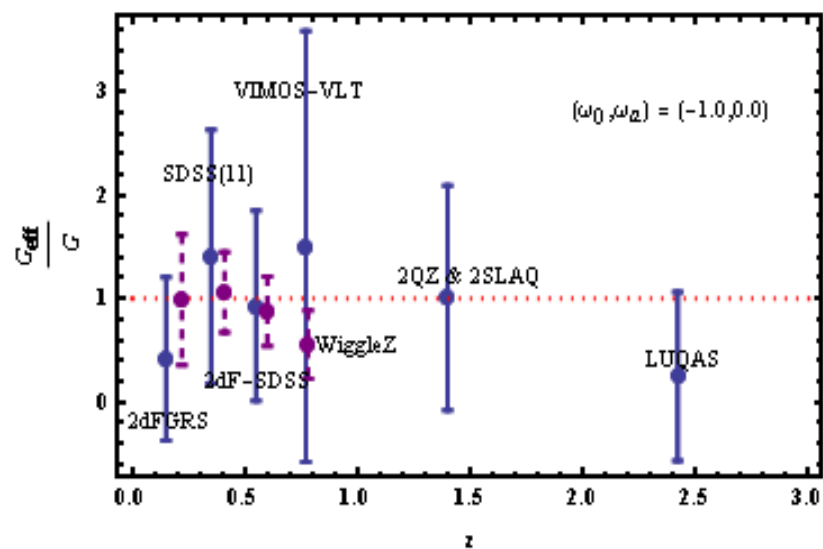
M : absolute magnitude @ $z \sim 0.5$

Matter Growth

Observations

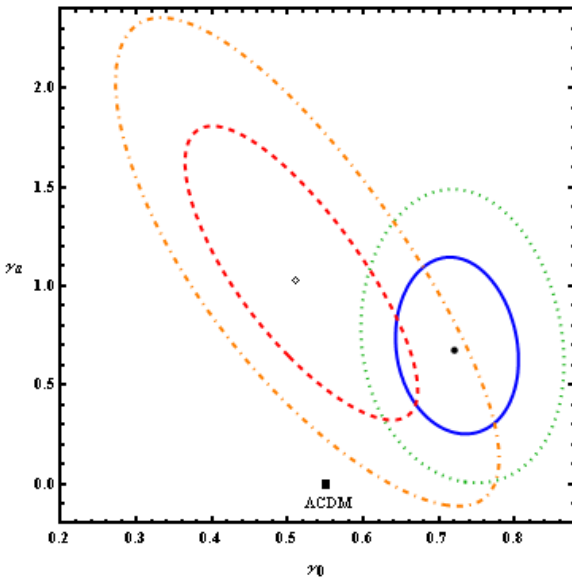
TABLE I: Data of the growth rate of clustering. The correspondence of the columns is as follows: number, redshift, observed growth rate, cosmological parameters used by different authors and references.

Index	z	$f_{obs,Ref}$	$(\Omega_{m,Ref}, \sigma_{8,Ref})$	Refs.
1	0.15	0.49 ± 0.14	(0.30, 0.90)	[16, 40, 41]
2	0.35	0.70 ± 0.18	(0.24, 0.76)	[42]
3	0.55	0.75 ± 0.18	(0.30, 1.00)	[43]
4	0.77	0.91 ± 0.36	(0.27, 0.78)	[16]
5	1.40	0.90 ± 0.24	(0.25, 0.84)	[44]
6	2.42	0.74 ± 0.24	(0.26, 0.93)	[19, 45]
7	3.00	1.46 ± 0.29	(0.30, 0.85)	[46]
8	0.22	0.60 ± 0.10	(0.27, 0.80)	[38]
9	0.41	0.70 ± 0.07	(0.27, 0.80)	[38]
10	0.60	0.73 ± 0.07	(0.27, 0.80)	[38]
11	0.78	0.70 ± 0.08	(0.27, 0.80)	[38]

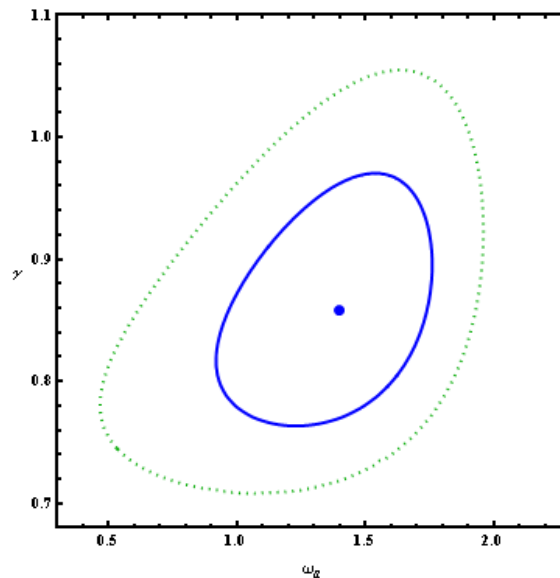


Matter Growth Constraints

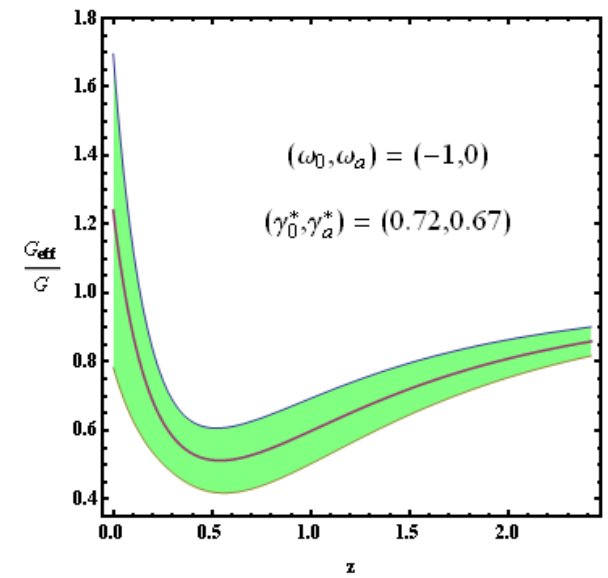
1



2



3



From current observations

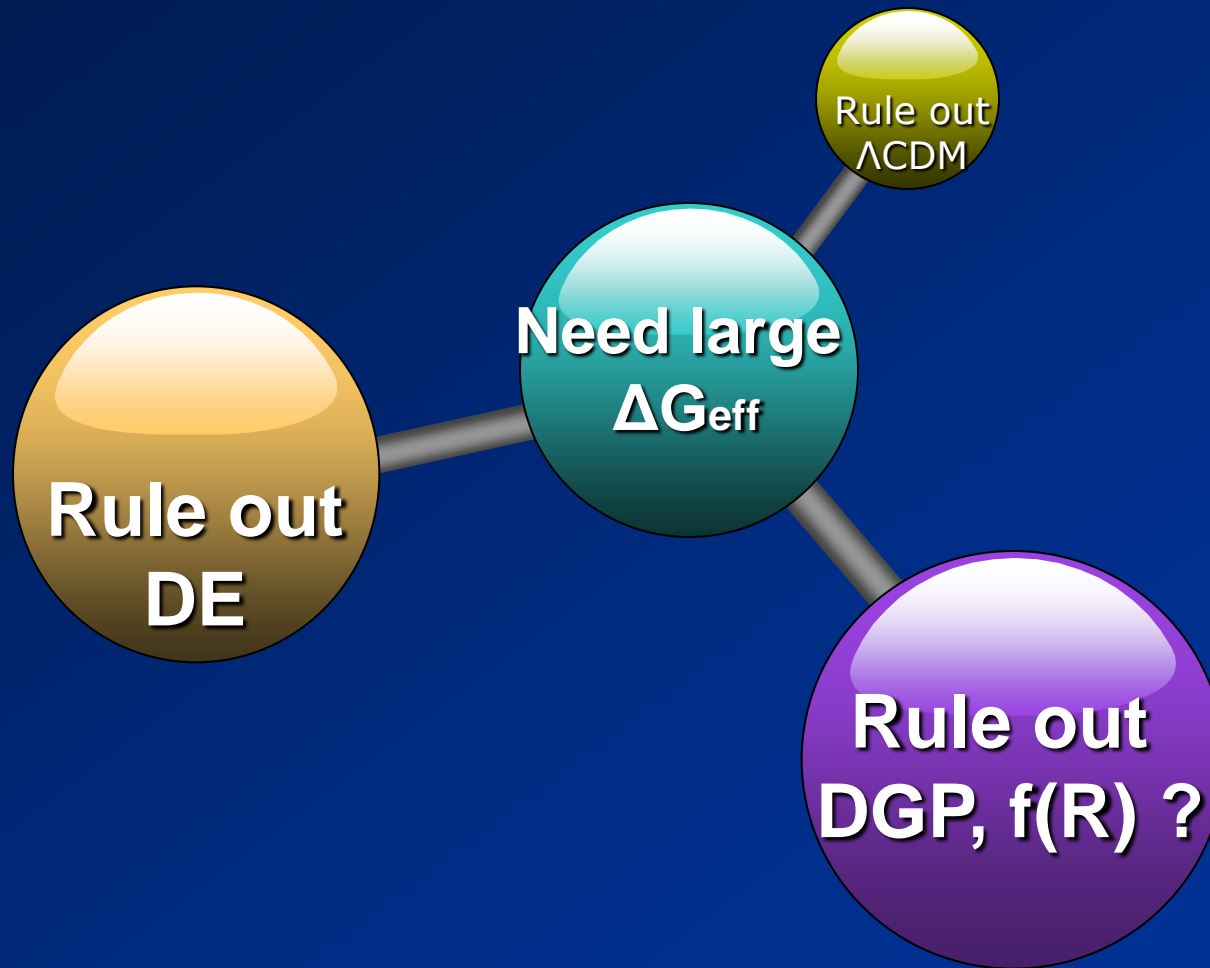
For Λ CDM, $\gamma \sim 0.56$

For DGP, $w_0 = -0.78, w_a = 0.32, \gamma \sim 0.68$

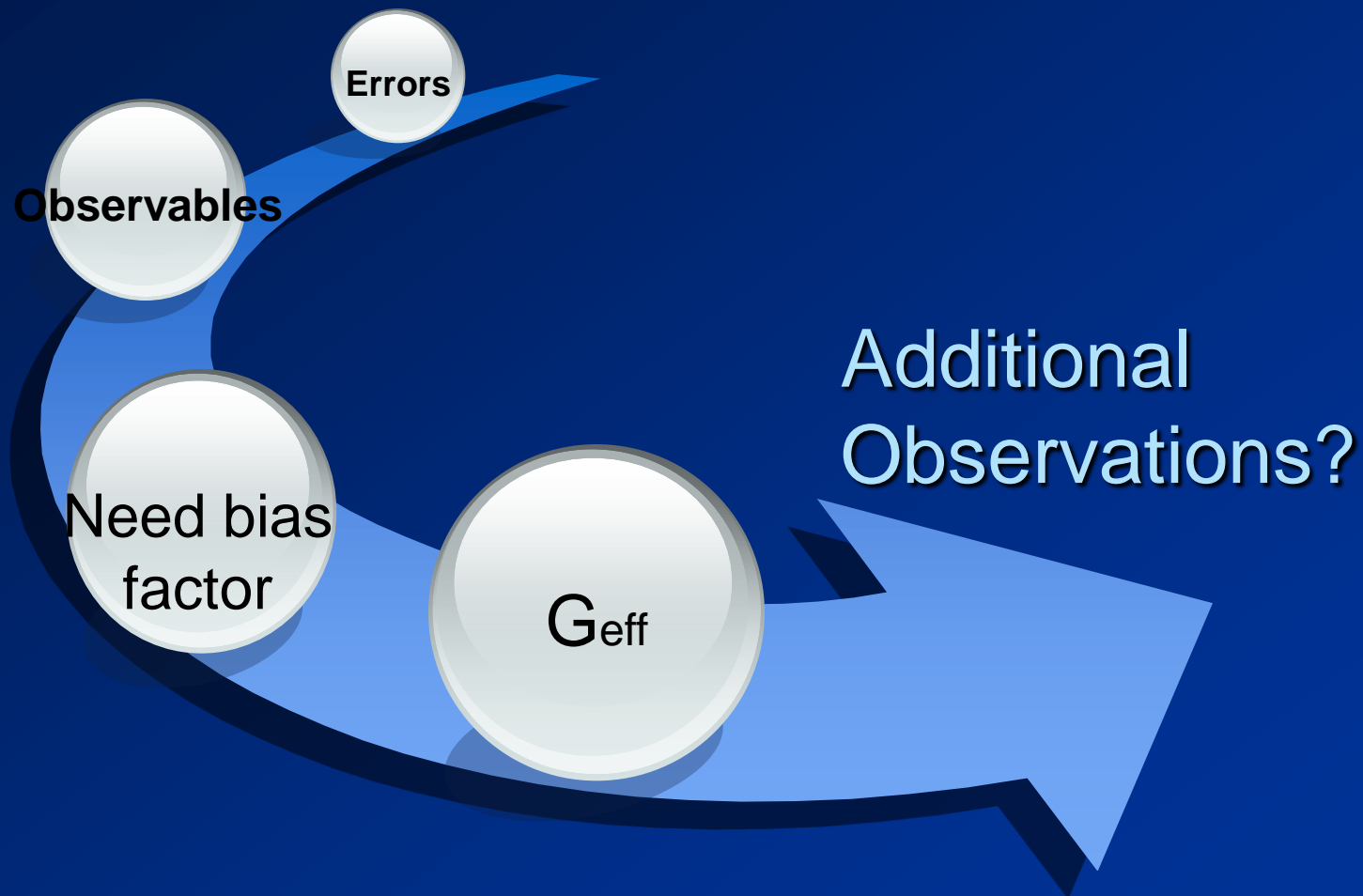
For STG, we may(?) produce large γ

Most of the
known models
rules out by $2\text{-}\sigma$

Conclusion



Challenges in observations



Zero Point Energy

(Emerging science, 1948...)

What?

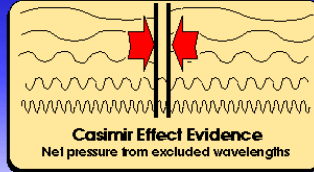
- Random Electromagnetic waves remain after all energy is removed
- Enormous energy density: 10^{24} to 10^{58} Joules/m³
- Theorized to indirectly cause gravity and inertia

Why?

- As an energy source?
- As a reactive medium?

Evidence?

- Casimir Effect
- Plank blackbody spectrum
- quantum effects



C.D. 94-0007

Thank You !

感謝